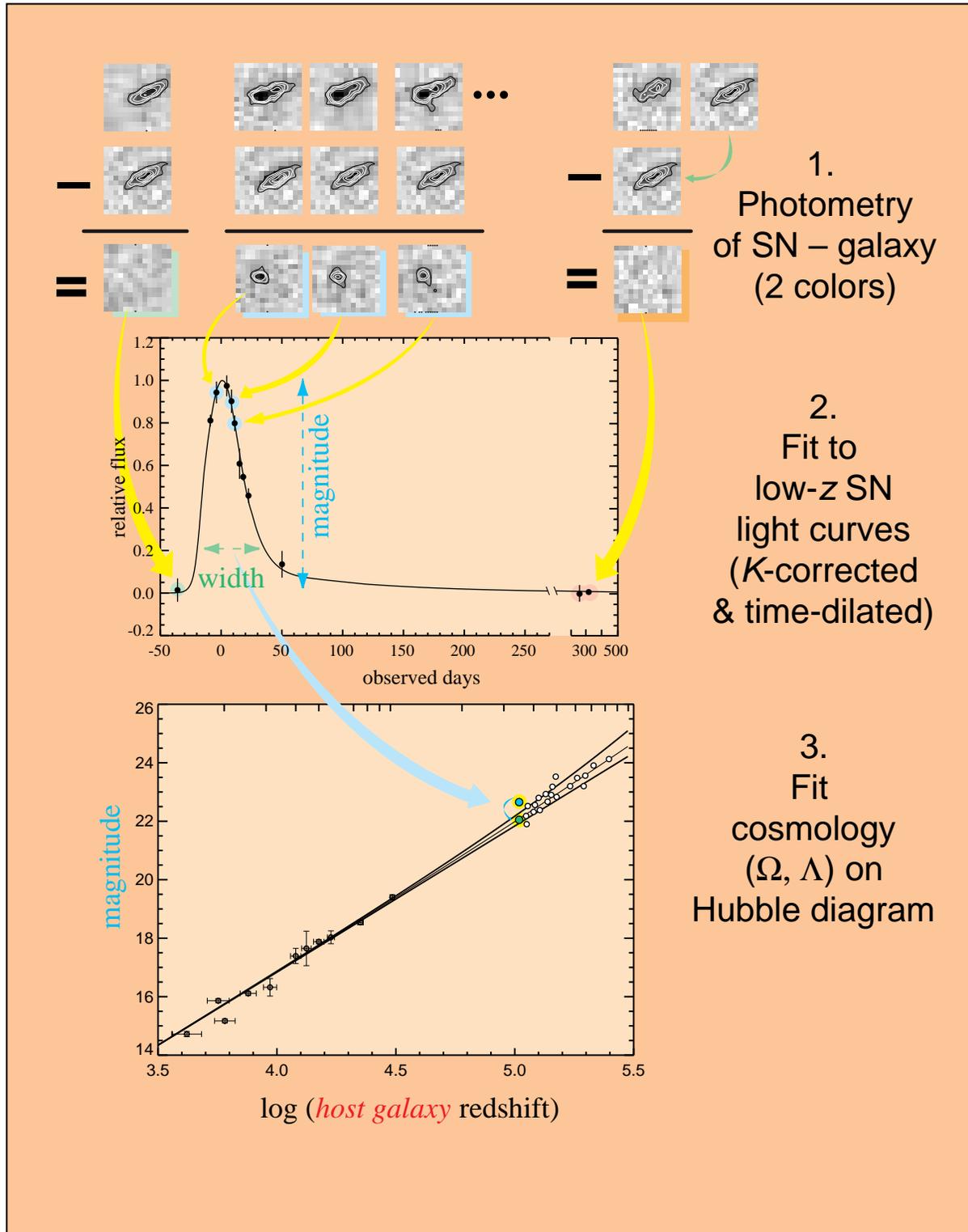


Analysis Steps



The supernovae are analyzed in the following three steps: First, the final image of the host galaxy alone is subtracted from the many images of each supernova spanning its lightcurve. The resulting *R*- and *I*-band photometry points are then fit to *K*-corrected (see Kim, Goobar, & Perlmutter, *P.A.S.P.* 1996) and $(1+z)$ -time-dilated *B*- and *V*-band template SN Ia lightcurves. This fit yields the apparent magnitude at peak and the best fit “stretch factor” that indicates the timescale (and hence the intrinsic luminosity) of each supernova. Finally, all of the supernova magnitudes—corrected for the stretch-luminosity relation—are plotted on the Hubble diagram as a function of their host galaxy redshift (when available, or supernova redshift, when not). The magnitudes vs. redshifts can then be fit to various alternative cosmologies. We fit the two “favorite” one-dimensional cases, the flat ($\Omega_M + \Omega_\Lambda = 1$) universe, and the $\Lambda = 0$ universe, as well as solving for a confidence region in the Ω_M -vs- Ω_Λ plane.